

Rapid and multi-band variability of the TeV-bright active nucleus of the galaxy IC 310

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ABSTRACT

Context. The galaxy IC 310 has recently been identified as a γ -ray emitter by the *Fermi*-LAT and at very high energies (VHE, $E > 100$ GeV) by the MAGIC telescopes. Originally classified as a head-tail radio galaxy, the nature of this object is subject of controversy because its nucleus shows blazar-like behavior.

Aims. In order to understand the nature of IC 310 and the origin of the VHE emission we studied the spectral and flux variability of IC 310 from the X-ray band to the VHE γ -ray regime.

Methods. The daily light curve of IC 310 above 300 GeV has been measured with the MAGIC telescopes from 2009 October to 2010 February. Contemporaneous *Fermi*-LAT data (2008–2011) in the 10–500 GeV energy range were also analyzed. In the X-ray regime, archival observations from 2003 to 2007 with *XMM-Newton*, *Chandra*, and *Swift*-XRT in the 0.5–10 keV band were studied.

Results. Several flares with similar amplitude can be seen in the MAGIC light curve. Day to day flux variability is clearly present ($> 5\sigma$). The photon index between 120 GeV and 8 TeV is measured to be $\Gamma \sim 2.0$ during both low and high flux states. It does not show variability, whereas the flux at 1 TeV changes by a factor of ~ 7 . *Fermi*-LAT detected only eight γ rays in the energy range 10 GeV–500 GeV in three years of observation. The measured photon index of $\Gamma = 1.3 \pm 0.5$ in the *Fermi*-LAT range is very hard. The X-ray measurements show strong variability in both flux and photon index. The latter varied from 1.76 ± 0.07 to 2.55 ± 0.07 .

Conclusions. The rapid variability measured in γ rays and X-rays confirms the blazar-like behavior of the source. The spectral behavior measured in both energy bands suggest IC 310 could be linked to extreme high frequency peaked BL Lac objects (HBL). The apparent luminosity of IC 310 is a few orders of magnitude lower than other extreme HBLs however and, atypical for an HBL, it harbors a kiloparsec-scale radio structure. Thus, IC 310 could be a peculiar type of a HBL, i.e., an intermediate state between BL Lac and FR I radio galaxy.

Key words. Galaxies: active – Galaxies: individual: IC 310 – Gamma rays: galaxies – X-rays: galaxies

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1. Introduction

The galaxy IC 310 (redshift of $z = 0.0189$; Bernardi et al. 2002) is one of the brightest objects of the Perseus cluster of galaxies

at radio frequencies as well as at X-ray energies. It has been detected in the high energy γ -ray band above 30 GeV with the *Fermi* Large Area Telescope (*Fermi*-LAT; Neronov et al. 2010) and above 260 GeV with the MAGIC telescopes (Aleksić et al. 2010).

IC 310 was originally classified as a head-tail radio galaxy (HTRG; Sijbring & de Bruyn 1998) based on the radial alignment of its radio jet on kiloparsec scale with the radially oriented pressure gradient in the surrounding intracluster medium (ICM). The radio morphology of such radio galaxies consists of a bright “head”, located at the core of the host galaxy, and “tails” pointing away from the center of the cluster. In the case of IC 310, the jet bending was assumed to be large, giving rise to the classification as a narrow-angle HTRG (Ryle & Windram 1968; Sijbring & de Bruyn 1998; Miley 1980; Lal & Rao et al. 2005; Feretti et al. 1998).

A recent detailed investigation of the radio structure of IC 310 using VLBA observations (Kadler et al. 2012) questions this classification. Inside the “head”, a parsec-scale core-jet structure was detected with no significant detection of a counter jet. The parsec-scale jet appears oriented in the same direction as the kiloparsec structure (the “tail”). This morphology disagrees with the classification as a HTRG since there is no indication for a jet bending process that determines the direction of the tail. Instead, the missing counter-jet implies that Doppler-boosting plays a significant role in this object. This behavior is naturally seen in blazars, i.e., AGN with their jets pointing along the line of sight. Thus, IC 310 appears to be more similar to other AGN detected in the very high energy regime (VHE, >100 GeV) that are mostly blazars (Ackermann et al. 2011).

If there is no bend, the ~ 400 kpc scale (projected) radio jet (Sijbring & de Bruyn 1998) would indicate a large angle between the jet and the line of sight. Therefore, IC 310 could belong to the same class as the three other “non-blazar” AGN discovered at VHE, M 87 (Aharonian et al. 2003; Aharonian et al. 2006; Acciari et al. 2008; Albert et al. 2008), Centaurus A (Aharonian et al. 2009), and NGC 1275 (Aleksić et al. 2012a) which are classified as Fanaroff-Riley I (FR I) radio galaxies. Hence, IC 310 could be either the closest TeV blazar or the brightest radio galaxy at TeV energies. IC 310 could belong to a new intermediate-state class in the AGN unification scheme, in between FR I radio galaxies and BL Lac objects.

The classification of IC 310 in a transitional population between BL Lac and FR I radio galaxy was already suggested by Rector et al. (1999), based on the optical, radio and X-ray properties of the source. Its weak optical emission lines are similar to those typically found in FR I radio galaxies but the non-thermal continuum from radio to the X-ray range is comparable to a low-luminosity BL Lac (Owen et al. 1996). In X-rays, the object is strongly dominated by the non-thermal point-like emission coincident with the radio “head” (Schwarz et al. 1992; Rhee et al. 1994; Sato et al. 2005). A faint X-ray halo has been resolved by *Chandra* observations extending in the direction of the radio tail as well (Dunn et al. 2010).

Aleksić et al. (2010) reported the detection of IC 310 during observations of the Perseus cluster with MAGIC taken place between 2008 November and 2010 February. A hard spectrum ($F \propto E^{-\Gamma}$ with $\Gamma = 2.0 \pm 0.14$) between 150 GeV and 7 TeV and indications for flux variability on time scales of months and years were reported. The variability time scales of years is confirmed by the non-detection of the source reported in 2010 August - 2011 February (Aleksić et al. 2012a) as well as by a re-detection in more recent MAGIC data (ATel #4583).

We present a re-analysis of the MAGIC stereo-observation taken during the period of strong activity between 2009 October and 2010 February allowing faster sampling of the light curve and more accurate spectral analysis. We also investigate γ -ray data of the *Fermi*-LAT instrument between 2008–2011 and archival X-ray observations with *XMM-Newton* (2003), *Chandra* (2004 and 2005) and *Swift*-XRT (2007). In Sect. 2, the γ -ray and X-ray observations and methods of data extraction and analysis are presented, while the resulting light curves and energy spectra are reported in Sect. 3. Then, interpretations of these results are discussed in Sect. 4.

2. Data: observations and analysis

2.1. VHE γ -ray data: The MAGIC telescopes

VHE observations of IC 310 were carried out with the MAGIC telescopes which are two Imaging Atmospheric Cherenkov Telescopes located on the island of La Palma at an altitude of 2200 m. Both telescopes consist of a mirror dish of 17 m diameter associated with a fast imaging camera of 3.5° field of view. The trigger threshold is ~ 50 GeV and the sensitivity above 290 GeV (in 50 h) is $\sim 0.8\%$ of the Crab Nebula flux with an angular resolution better than 0.07° (Aleksić et al. 2012b).

From 2009 October to 2010 February, the Perseus cluster was observed with MAGIC in the so-called wobble mode, i.e. pointing alternatively to two positions 0.4° away from the center of the cluster (NGC 1275). Since IC 310 is located 0.6° from the cluster center, it appeared as an off-axis source in these observations. The position of the source is 0.25° and 1° away from the camera center in the two wobble positions, respectively. In our previous publication on IC 310 (Aleksić et al. 2010) only data of the closest wobble position were used to construct the light curve and the spectrum. Here, we used an improved analysis method that can handle the data of both wobble positions and hence provides more accurate results.

After data quality selection mainly based on the atmospheric conditions, the data sample corresponds to 43.3 h effective time, t_{eff} . The calibration, image cleaning, parameterization, and event reconstruction as well as the gamma/hadron separation were performed with the standard analysis software MARS described in Moralejo et al. (2009). The background estimation is done separately for each wobble position. For the closer wobble position 2 OFF (signal-free) regions at 0.25° away from the camera center were used and for the wobble position further away 5 OFF regions were chosen with an offset of 1° , respectively. The γ -ray signal is calculated by subtracting the estimated background from the on-source region events. The effective area, A_{eff} , of MAGIC strongly depends on the distance from the camera center. It is estimated separately for each wobble position using Monte Carlo simulations of γ rays initiated at the same distance from the camera center as the source position.

The spectrum and the light curve from both wobble positions are combined by calculating the sum, over two wobble positions, of the total number of excess events, N_{ex} , measured during each wobble observation weighted by the corresponding effective area and effective time.

Because IC 310 is not located in the central position, the systematic errors can be higher than reported for the standard 0.4° wobble observation by Aleksić et al. (2012b). In order to study the systematic effects, we analyzed Crab Nebula observations taken at different offsets from the camera center (from 0.2° to 1.4°) with the same analysis chain. The Crab spectra measured at different offsets, e.g. with 0.2° and 1° , are in good agreement

with the standard observation result. The systematic errors on the flux normalization and photon index are estimated to be below $\sim 17\%$ and ~ 0.2 , respectively (instead of 11% and 0.15 for standard wobble observations; Aleksić et al. 2012b). The systematic uncertainty on the energy scale is evaluated to be 15% (Aleksić et al. 2012b).

2.2. HE γ -ray data: Fermi-LAT

The *Fermi*-LAT is a pair-conversion telescope sensitive to photons between 20 MeV and several hundred GeV (Atwood et al. 2009; Ackermann et al. 2012). Since August 5, 2008 it has operated primarily in sky survey mode, scanning the entire sky every three hours. The data used in this paper were taken between August 5, 2008 and July 31, 2011 (MJD 54683–55773), overlapping the MAGIC observations. The *Fermi*-LAT results presented here were obtained with the analysis pipeline used to produce the “The First Fermi-LAT Catalog of Sources Above 10 GeV”, designated 1FHL, which is currently under preparation within the LAT Collaboration. This analysis was performed with the ScienceTools software package version v9r26p02. The results presented here were obtained with “Clean” class events in the energy range 10–500 GeV from the region centered at (R.A., Dec.) = (52:921, 41:634) (J2000), and has a radius of 6.198° . Moreover, only data for time periods when the spacecraft rocking angle was less than 52° were used, and events with zenith angles larger than 105° were excluded in order to reduce the contamination from Earth limb γ -rays, which are produced by cosmic rays interacting with the upper atmosphere.

In the analysis, the Galactic and extragalactic diffuse backgrounds were parameterized with the files `gal_2yearp7v6_v0.fits` and `iso_p7v6clean.txt`, which are publicly available¹. Because of the relatively small size of the region, the limited photon count, and the relatively small effective energy range (most photons cluster in the energy range 10–100 GeV), there is some degeneracy in the simultaneous characterization of the Galactic and isotropic diffuse components, so we fixed the normalization of the isotropic component to the best-fit value over the entire sky, and left free the normalization of the Galactic component.

The region analyzed contains only two sources detected above >10 GeV, which based on close positional agreement we have associated with IC 310 and NGC 1275. The source positions used in the spectral fit were the optimized positions from the LAT analysis, which are (49:169, 41:322) for IC 310 and (49:977, 41:501) for NGC 1275. These positions differ from the actual source positions by $0:01$ and $0:02$ respectively for IC 310 and NGC 1275, which are well within the 95% confidence regions (error ellipses) for these two objects: $0:09$ and $0:03$. The source associated with NGC 1275 is ~ 20 times brighter than IC 310 at GeV energies and separated by only $\sim 0:65$. We note that above 10 GeV the 68% containment radius of the point-spread function of the LAT is about $0:2$, which means that the sources are well resolved, and in particular the spectrum of IC 310 can be separately measured.

We used simple power-law models to characterize the spectra of the sources. The spectral fitting was performed with the binned likelihood method using the P7_V6_CLEAN instrument response functions (see Ackermann et al. 2009), ten bins per decade in energy starting at 10 GeV, and an angular binning of $0:05$ and $0:1$ for γ -rays that converted in the thin and thick

tungsten layers of the tracker, respectively. Following the global fitting over the full energy range, we extracted photon fluxes in three energy bands: 10–30 GeV, 30–100 GeV and 100–500 GeV by fixing the photon indices (to those from the overall spectral fit) and leaving free only the normalizations in the fits. Because of the low statistics, the flux errors are strongly dominated by Poisson fluctuations and so are not symmetric. In this manuscript we report separate 1σ uncertainties toward low/high fluxes obtained via MINOS in the MINUIT package.

2.3. X-ray data: Chandra, XMM-Newton, Swift-XRT

Two *Chandra* observations of IC 310 were taken with the Advanced CCD Imaging Spectrometer (ACIS) in the 0.5–8 keV band (ACIS-I Observation ID 5596 and 5597). The data sets are not affected by pile-up or detector heating. Both observations were made within four months: December 26, 2004 with effectively 25.2 ks for Obs. ID 5597 and March 23, 2005 with 1.5 ks for Obs. ID 5596. The data were analyzed with the Chandra Interactive Analysis of Observations (CIAO 4.4) software using Version 4.4.8 of the calibration files. The extraction radii were chosen to be $4.92''$ for Obs. ID 5597 and $2.46''$ for Obs. ID 5596.

One observation of IC 310 with an exposure of 22.6 ks was taken on February 26, 2003 (Obs. ID 0151560101) with *XMM-Newton*. Here, we analyzed only data taken with the pn detector of the European Photon Imaging Camera (EPIC-pn) covering the energy range 0.2–15 keV (Strüder et al. 2001). Data were reduced using the *XMM-Newton* Software Analysis System (SAS v.11.0.0) and the newest calibration files. For the source spectra, we used a circular region of $30''$ radius, centered on the source. The background was extracted from an equally large circle, but located outside any distinct source of radiation. Only single and double events were used to derive the spectrum.

Swift consists of several instruments covering a wide range of wavelengths including the X-ray telescope (XRT; Burrows et al. 2005) operating in the 0.2–10 keV band. The observation of IC 310 was performed on February 19, 2007 with an effective observation time of 4.1 ksec. The data were reduced with XRTPIPELINE (V 0.12.6). For the source region, a circle with a radius of $47''$ was used. The background region was created using an annulus centered on the source coordinates with an inner radius of $118''$ and an outer radius of $165''$. The extraction of events was done in XSELECT (V2.4b), using event grades 0–12.

The X-ray data analysis was performed with the *Interactive Spectral Interpretation System* (Houck & Denicola 2000). In order to allow for a reasonable spectral fitting, we re-binned the data according to a signal-to-noise criteria for each bin (eight for *Swift* and *XMM-Newton*, and four in the case of the two *Chandra* observations). For the spectral fitting, we used data between 0.5 keV and 10 keV for all three instruments. The photon index Γ and the column density of the neutral absorption was determined by the spectral fitting. As IC 310 is located $0:6$ away from the center of the cluster, the contribution of the thermal ICM emission to the X-ray signal is very small, and therefore, it was not necessary to include it when modeling the data.

3. Results

3.1. VHE lightcurve

The integral VHE γ -ray flux of IC 310 has been derived assuming a differential spectrum for the calculation of the effective area following a power-law with a photon index $\Gamma = -2.0$ as

¹ See <http://fermi.gsfc.nasa.gov/ssc/data/access/lat/BackgroundModels.html>

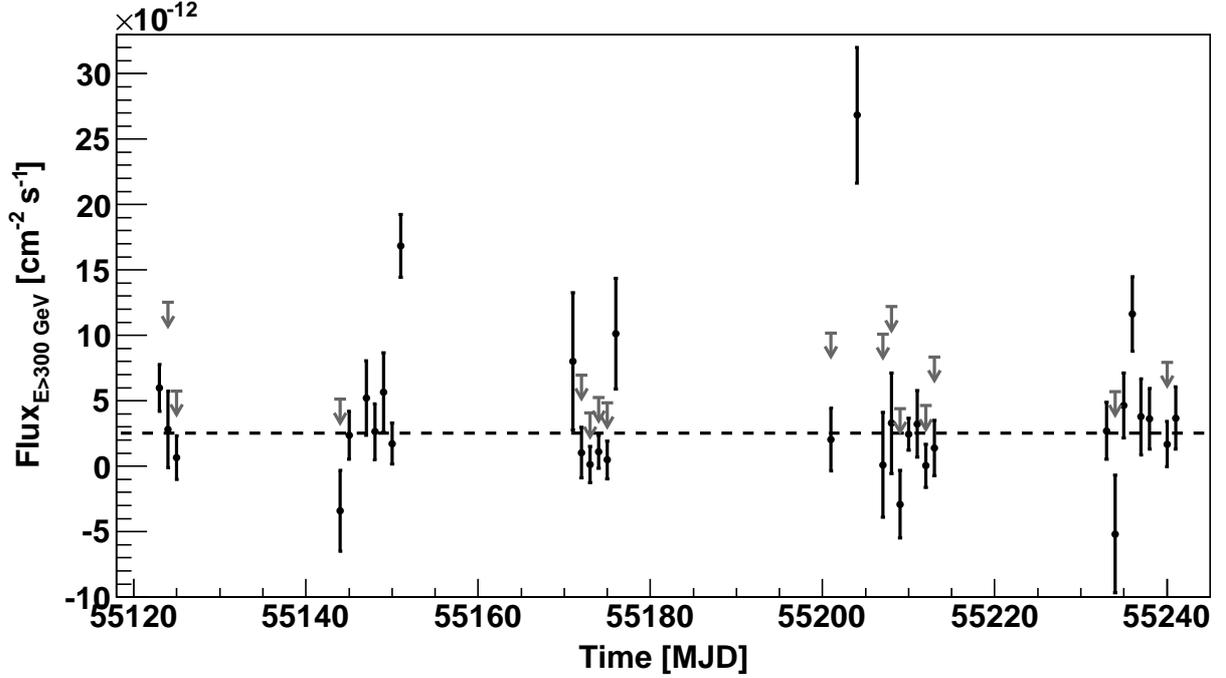


Fig. 1. Daily light curve of IC 310 above 300 GeV from 2009 October to 2010 February. The arrows are 95% confidence level upper limits calculated for days compatible with no signal. The dashed line shows the constant flux fit to all flux points not taking into account the upper limits ($\chi^2/\text{d.o.f} = 102/32$).

found in Aleksić et al. (2010). However, the dependence of the resulting light curve on the index is only minor. The flux above 300 GeV measured with MAGIC during 33 individual days between 2009 October and 2010 February is shown Fig. 1 and listed in Table 1. Integral upper limits have also been calculated for days showing an excess below one standard deviation (σ). The upper limits are determined by applying model 4 of Rolke et al. (2005), using a confidence level (c.l.) of 95% and 30% systematic uncertainty. The mean flux above 300 GeV during this period is $\Phi_{\text{mean}} = (3.62 \pm 0.40) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$ which is in good agreement with the mean flux reported previously, $\Phi_{\text{mean}} = (3.1 \pm 0.5) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$, using only the 0°25 wobble position (Aleksić et al. 2010). The best fit of the flux points, i.e. not including the upper limits, with a constant flux (dashed line, $\Phi_{\text{CstFit}} = (2.52 \pm 0.37) \times 10^{-12} \text{ cm}^{-2} \text{ s}^{-1}$) has a χ^2 test value of 102 for 32 degrees of freedom corresponding to a probability of 3.10×10^{-9} that the source is not variable. Note, that since the high χ^2 value rejects the validity of the constant flux fit, this value cannot be interpreted as the average flux of the source in the observed time span. The variability of the source is thus detected with a confidence level of 5.8σ .

Three days show a flux $> 3\sigma$ above the constant flux fit: November 16, 2009 (MJD = 55151), January 8 (MJD = 55204) and February 9, 2010 (MJD = 55236)². The mean flux during these three days is $(1.60 \pm 0.17) \times 10^{-11} \text{ cm}^{-2} \text{ s}^{-1}$, which is more than six times higher than the constant fit flux. Excluding these three days, the light curve is compatible with a constant flux (probability of 36%). The three flares are isolated and seem to be restricted to a single day bin. The most significant one happened in 2009 November. Figure 2 shows a zoom-in of the daily light curve for this month. IC 310 was observed every day from 9 to 16 except for November 11. The probability of a constant flux

² Observations were taken around midnight UTC. The dates given in this paper correspond to the following day. Exact observation times are given in Table 1.

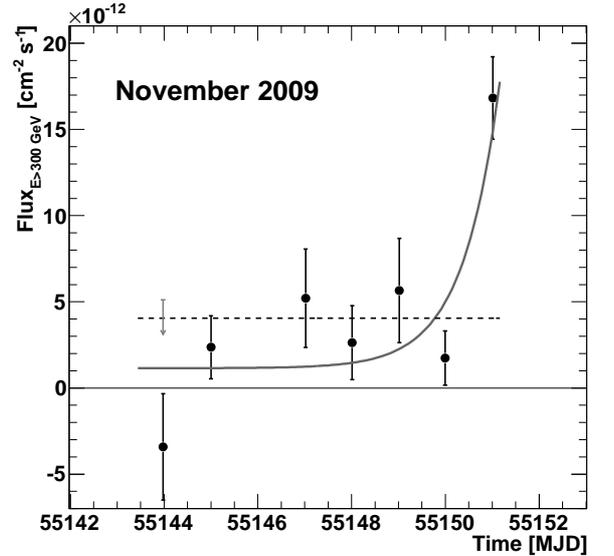


Fig. 2. Daily light curve of IC 310 above 300 GeV in 2009 November. The black dashed line shows the best constant fit of the data. The thick grey line shows the constant plus an exponential component corresponding to the doubling time of $\tau_{\text{Nov,UL}} = 0.55 \text{ d}$ providing an acceptable fit of the data (5% of probability).

during this period is 10^{-6} . The flare appeared during the last day of observation in November and no evidence of an increased flux is seen before. The variability time scale might be shorter than the daily-scale sampling of the light curve but, due to limited statistics intra-night variability could not be established.

The characteristic flux doubling time during a flare can be estimated by fitting the data with a constant plus an exponential

Table 1. Results from individual days of MAGIC observations.

used data ^a	MJD start	t_{eff} [h]	$F_{E>300\text{GeV}}^b$ [$10^{-12}\text{ ph cm}^{-2}\text{ s}^{-1}$]
all data		43.32	3.62 ± 0.40
2009-10-19	55123.02	2.43	5.97 ± 1.79
2009-10-20	55124.05	1.11	$(2.80 \pm 2.93) < 12.52$
2009-10-21	55125.02	1.58	$(0.67 \pm 1.67) < 5.74$
2009-11-09	55143.96	0.80	$(-3.42 \pm 3.09) < 5.11$
2009-11-10	55144.97	1.19	2.36 ± 1.83
2009-11-12	55146.98	0.99	5.21 ± 2.86
2009-11-13	55147.98	1.17	2.64 ± 2.14
2009-11-14	55148.98	1.11	5.65 ± 3.03
2009-11-15	55149.94	2.61	1.73 ± 1.57
2009-11-16	55150.94	2.34	16.83 ± 2.40
2009-12-06	55170.89	0.44	8.01 ± 5.25
2009-12-07	55171.89	1.44	$(1.03 \pm 1.94) < 6.95$
2009-12-08	55172.88	2.51	$(0.13 \pm 1.39) < 4.06$
2009-12-09	55173.89	3.18	$(1.10 \pm 1.25) < 5.25$
2009-12-10	55174.88	2.40	$(0.49 \pm 1.44) < 4.83$
2009-12-11	55175.88	0.54	10.13 ± 4.23
2010-01-05	55200.83	0.81	$(2.05 \pm 2.40) < 10.02$
2010-01-08	55203.89	0.78	26.84 ± 5.19
2010-01-11	55206.87	0.62	$(0.09 \pm 4.00) < 10.09$
2010-01-12	55207.90	0.61	$(3.29 \pm 3.85) < 12.19$
2010-01-13	55208.88	0.86	$(-2.91 \pm 2.59) < 4.42$
2010-01-14	55209.84	1.47	2.44 ± 1.22
2010-01-15	55210.83	1.46	3.23 ± 2.54
2010-01-16	55211.83	1.90	$(0.03 \pm 1.66) < 4.63$
2010-01-17	55212.84	1.22	$(1.40 \pm 2.11) < 8.33$
2010-02-06	55232.84	1.22	2.71 ± 2.18
2010-02-07	55233.87	0.60	$(-5.19 \pm 4.50) < 5.71$
2010-02-08	55234.85	0.90	4.65 ± 2.47
2010-02-09	55235.85	1.38	11.65 ± 2.84
2010-02-10	55236.87	0.62	3.77 ± 2.91
2010-02-11	55237.86	0.87	3.62 ± 2.32
2010-02-13	55239.84	0.92	$(1.69 \pm 1.74) < 7.95$
2010-02-14	55240.84	1.24	3.68 ± 2.38

Notes. ^(a) Dates in MAGIC night notation. ^(b) Measured flux above 300 GeV in units of $10^{-12}\text{ ph cm}^{-2}\text{ s}^{-1}$. Upper limits are given with 95% confidence level.

increase. However, in the night just before the flare in November the flux was particularly low and the best fit doubling time goes to zero. In order to estimate the doubling time that allows us to measure such variability, we fit the data with fixed doubling times and calculate the probability of each hypothesis with the χ^2 method. The doubling time providing a fit probability above 5% is $\tau_{\text{Nov,UL}}=0.55\text{ d}$ (thick line in Fig. 2).

The 2010 January and February flares provide much less constraining results because of a sparse time coverage in January and a much less significant flare in February. For the rest of the paper, we consider only the upper limit obtained from the November data.

3.2. VHE spectra

For the spectral analysis, we split the MAGIC data set into two samples corresponding to different γ -ray emission states. We define a “high state” containing the three days with a higher flux level, $F_{E>300\text{GeV}} > 1.1 \times 10^{-11}\text{ cm}^{-2}\text{ s}^{-1}$, (MJD 55151, 55204 and 55236, $t_{\text{eff}}=4.5\text{ h}$) and a “low state” including all the other days ($t_{\text{eff}}=38.8\text{ h}$). The reconstructed spectra between 120 GeV and 8.1 TeV for both states are shown in Fig. 3. Due to the absorption of VHE γ rays by the Extragalactic Background Light

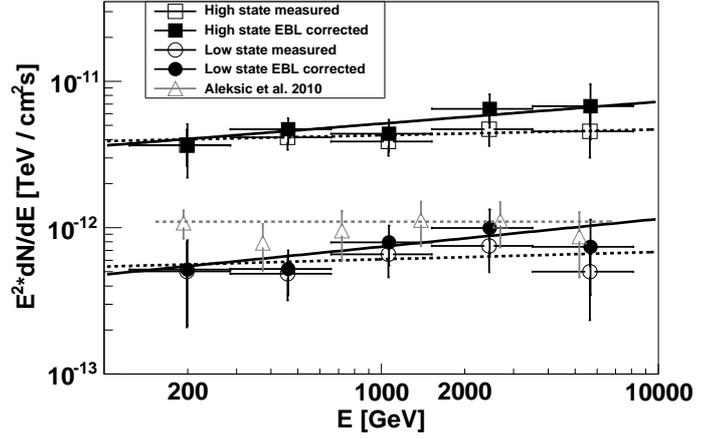


Fig. 3. Measured (open markers) and EBL-absorption corrected (filled markers) spectral energy distribution for IC 310 obtained by MAGIC in high and low states together with their power-law fits in dashed line (measured) and solid line (EBL corrected). Fit parameters are given in Table 1. For comparison, we show the result reported in Aleksić et al. 2010 (grey triangles, without EBL de-absorption) for the whole period using only the closest pointing observations.

Table 2. Results of power-law fit of the 0.12–8.1 TeV spectra measured with MAGIC.

state		$f_0 \pm f_{\text{stat}} \pm f_{\text{syst}}$ $\times 10^{-12} [\text{TeV}^{-1}\text{ cm}^{-2}\text{ s}^{-1}]$	$\Gamma \pm \Gamma_{\text{stat}} \pm \Gamma_{\text{syst}}$
high	observed	$4.28 \pm 0.21 \pm 0.73$	$1.96 \pm 0.10 \pm 0.20$
	intrinsic	$5.14 \pm 0.28 \pm 0.90$	$1.85 \pm 0.11 \pm 0.20$
low	observed	$0.608 \pm 0.037 \pm 0.11$	$1.95 \pm 0.12 \pm 0.20$
	intrinsic	$0.741 \pm 0.045 \pm 0.14$	$1.81 \pm 0.13 \pm 0.20$

(EBL) through pair creation, the observed spectra of extragalactic sources are softened. We corrected the spectra for the EBL absorption according to different EBL models (Dominguez et al. 2011; Franceschini et al. 2008; Kneiske & Dole 2010) which all provide comparable results within the systematic errors. The EBL correction applied here uses the model by Dominguez et al. (2011). Due to the proximity of IC 310 ($z = 0.0189$), the effect of the absorption is relatively modest, reducing the flux above 1 TeV by 15–20% and steepening the photon index by ~ 0.1 .

Both observed and de-absorbed spectra can be well described by a simple power law for both emission states:

$$\frac{dF}{dE} = f_0 \times \left(\frac{E}{1\text{TeV}} \right)^{-\Gamma} \left[\frac{10^{-12}}{\text{cm}^2\text{s TeV}} \right]. \quad (1)$$

Results for the flux normalization at 1 TeV, f_0 , and the photon index, Γ , are summarized in Table 2. Changes in f_0 by a factor of ~ 7 between the low and high states had been observed, but no significant change in the photon index.

3.3. Fermi-LAT results

The three year data taken with *Fermi*-LAT resulted in the detection of IC 310 above 10 GeV with a test statistic of $\text{TS} = 27.0$ (4.5σ). The integrated flux F_{10-500} between 10 and 500 GeV is measured to be $(6.9 \pm 3.3) \times 10^{-11}\text{ cm}^{-2}\text{ s}^{-1}$. The spectrum in this

Table 3. Arrival timestamps and energies of γ -candidates above 10 GeV (from the first three years of *Fermi*-LAT accumulated data) from a 0.3 radius circle centered at the position of IC 310. The energy resolution of the *Fermi*-LAT instrument at the energies reported for these events is about 10% (Ackermann et al. 2012).

MJD	Energy [GeV]
54720.03	96.4
54833.95	112.1
54846.64	22.2
54972.38	12.6
55081.11	39.0
55118.56	148.3
55247.01	12.1
55462.98	46.3

energy band can be fitted by the following power-law formula with a very hard photon index of $\Gamma = 1.3 \pm 0.5$:

$$\frac{dF}{dE} = \frac{F_{10-500}(-\Gamma + 1)E^{-\Gamma}}{E_{\max}^{-\Gamma+1} - E_{\min}^{-\Gamma+1}} \left[\frac{1}{\text{cm}^2 \text{s GeV}} \right]. \quad (2)$$

Here, E_{\min} and E_{\max} are the lower and the upper boundary of the energy bins in GeV, respectively.

The estimated fluxes in three energy bands (10–30 GeV, 30–100 GeV and 100–500 GeV) are shown in Fig. 5. Because of the low statistics, the flux errors are strongly dominated by Poisson fluctuations. Thus, they are asymmetric. In this manuscript we report separate 1σ uncertainties toward low/high fluxes obtained via MINOS in the MINUIT package.

In *Fermi*-LAT data, seven of the eight photons detected above 10 GeV arrived within the first 1.5 years of observation (see arrival times in Table 3). This also suggests variability even if the low statistics do not allow a certain conclusion. No significant correlation between the *Fermi*-LAT photon arrival times and the MAGIC flares has been found.

3.4. X-ray behavior

Table 4 and Fig. 4 summarize the analysis results of archival X-ray data for an *XMM-Newton*, two *Chandra* and a *Swift*-XRT observations. The photon index, Γ , and the absorption column, N_{H} , are derived from fits of power-law models to the 0.5–10 keV data. Because intrinsic absorption due to material close to the X-ray source in excess of the neutral Galactic absorption towards IC 310 ($N_{\text{H}} = 0.12 \times 10^{22} \text{ cm}^{-2}$; Kalberla et al., 2010) cannot be excluded, N_{H} is left as a free parameter in these fits. Note that the obtained N_{H} values are compatible with the Galactic absorption except for *XMM-Newton*. We also tried an analysis of the *XMM-Newton* data with a broken power law and a fixed N_{H} . It resulted in a good $\chi^2/\text{d.o.f.}$ but the break energy was found to be at energies where the absorption takes place, i.e., at < 1 keV. Thus, a break in the spectrum is potentially mimicked by the absorption.

Our analysis shows that between 2003 and 2007 flux variability on time scales of years is present in the low energy regime (0.5–2 keV) as well as in the higher energy range (2–10 keV). This variability is accompanied by changes in the photon index Γ and N_{H} . A problem with this kind of analysis is that there is a well known correlation between both parameters, i.e., softer spectra with larger absorption columns cannot be distinguished from harder spectra with slightly smaller absorption. The right hand panel of Fig.4 shows confidence contours between Γ and N_{H} for all observations. While the *Swift* data have too low a

signal-to-noise ratio to make a statement on changes in the spectral parameters, a statistically very significant change in photon index and absorption column is present when comparing the higher quality *XMM-Newton* and *Chandra* data. The change in the spectral shape cannot be explained solely by a change in absorption, but the intrinsic spectrum of the source appears to have changed between the different observations.

We conclude that the intrinsic source spectral index varies between $\Gamma = 2.5$ (soft) and $\Gamma = 1.8$ (hard). During the *Chandra* observation the absorption column was consistent with the Galactic value towards IC 310, but N_{H} was significantly higher during the *XMM-Newton* observation, which could be due to an increase in the source intrinsic absorption. A possible cause for a change in the internal absorption of the source could be, e.g., the presence of material close to the black hole. Such variation has been seen in a few other active galaxies. For example, a long-term variation of the intrinsic absorption by 30% has been observed in Centaurus A (Benlloch et al. 2010), where recent observations also found evidence for a short term absorption event that lasted for ~ 170 d (Rivers et al. 2012).

4. Discussion

The rapid variability of the VHE γ -ray emission of IC 310 is firmly established by the day-scale light curve measured by the MAGIC telescopes. This rules out the previously discussed origin of the emission at a bow shock driven into the ICM by the radio jets (Neronov et al. 2010), as this should produce almost steady-state emission. Thus, the VHE γ rays more likely originate from the AGN.

The time variability allows us to constrain the size of the emission region. For a source located at a redshift, z , whose emission is beamed towards the observer with a corresponding Doppler factor, δ , the expression linking the radius of the emission region, R , with the variability timescale, τ , is:

$$R \leq c\tau_{\text{Nov,UL}}\delta(1+z)^{-1}. \quad (3)$$

Using the upper limit for the doubling time $\tau_{\text{Nov,UL}} = 0.55$ d and $z = 0.0189$ the radius is calculated to be $R \lesssim \delta \times 1.5 \times 10^{15}$ cm. For a more conservative estimate of the upper limit one can use the light curve binning of 1 day thus, $\tau = 1$ d as the observed flares seem to be limited to a single bin. This constrains the size of the emission region to be $R \lesssim \delta \times 2.5 \times 10^{15}$ cm. Comparing this with the Schwarzschild radius of the central black hole which is $R_{\text{S}} = 7 \times 10^{13}$ cm ($M_{\text{BH}} = 2.4 \times 10^8 M_{\odot}$ from the central velocity dispersion, Aleksić et al. 2010), the emission region must be very compact ($< 100 R_{\text{S}}$), assuming no Doppler boosting, and probably located in the inner-most jet. On the other hand, assuming a large Doppler factor, that emission could be as large as parsec scales, with the day-scale TeV variability explained by, e.g., the magnetic reconnection mini-jet model by Giannios (2012).

Figure 5 shows the spectral energy distribution (SED) extending from X-ray to VHE. Typically, νF_{ν} spectra of blazars feature two emission maxima, one in the optical/UV region and another in the MeV/GeV regime. The low energy hump is explained by synchrotron radiation of electrons assuming that the jet of an AGN consists of an electron-positron plasma (leptonic models, Hartman et al. 1992; Sikora 1994). By contrast, the origin of the high energy hump is still under discussion. In the mainstream interpretation, a one-zone Synchrotron-Self-Compton (SSC) scenario produces this peak by co-spatial Inverse Compton up-scattering of the synchrotron photons of

Table 4. Results of the analyzed X-ray observations.

instrument	date [MJD]	exposure [ks]	$F_{0.5-2\text{keV}}^a$ [$10^{-3} \text{ keV s}^{-1} \text{ cm}^{-2}$]	$F_{2-10\text{keV}}^b$ [$10^{-3} \text{ keV s}^{-1} \text{ cm}^{-2}$]	Γ^c	N_{H}^d [10^{22} cm^{-2}]	$\chi^2/\text{d.o.f.}$
<i>XMM-Newton</i>	52697	22.6	1.007 ± 0.012	$0.828^{+0.026}_{-0.040}$	$2.55^{+0.07}_{-0.04}$	$0.146^{+0.016}_{-0.008}$	124/104
<i>Chandra</i> ObsID 5596	53456	1.5	1.77 ± 0.13	2.5 ± 0.4	2.01 ± 0.20	$0.07^{+0.08}_{-0.07}$	62/78
<i>Chandra</i> ObsID 5597	53363	25.2	0.656 ± 0.019	1.39 ± 0.08	1.76 ± 0.07	$0.089^{+0.028}_{-0.027}$	97/78
<i>Swift-XRT</i>	54152	4.1	0.82 ± 0.10	$1.2^{+0.6}_{-0.5}$	$2.0^{+0.5}_{-0.4}$	$0.07^{+0.13}_{-0.07}$	12/16

Notes. ^(a) Measured flux between 0.5 and 2 keV determined by a simple power law fit. ^(b) Measured flux between 2 and 10 keV determined by a simple power law fit. ^(c) Photon index: $F \propto E^{-\Gamma}$. ^(d) Absorption with a equivalent column of hydrogen.

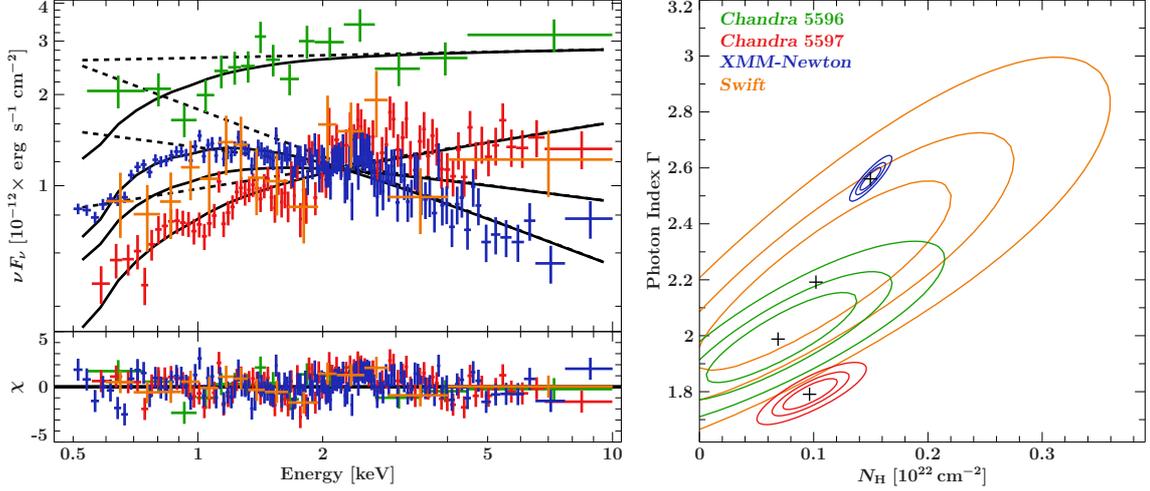


Fig. 4. *Left Top panel:* Measured spectral energy distribution of IC 310 in the X-ray band in different time periods. *Chandra* Obs. ID 5596 (green), *XMM-Newton* (blue), *Chandra* Obs. ID 5597 (red) and *Swift-XRT* (orange). The solid lines show the absorbed power law fit and the dashed lines show the de-absorbed fit, respectively. *Left Bottom panel:* Residuals of the χ^2 fit. *Right panel:* Contour plot between the intrinsic column density N_{H} and the X-ray photon index Γ for the spectra depicted in the left panel. The contours for each observation are given at 68%, 90%, and 99% confidence level. For comparison, the Galactic N_{H} value is $0.12 \times 10^{22} \text{ cm}^{-2}$.

the same energetic electron population which have just produced them. In hadronic models the origin of the high-energy hump can be explained by hadronic cascades, see e.g. Mannheim (1993) and Mücke et al. (2003). The multiwavelength characteristics of IC 310 presented here support the notion that the VHE emission is associated with a blazar-type jet but its apparent luminosity ($10^{42-43} \text{ erg s}^{-1}$) is two to three orders of magnitude lower than for typical TeV blazars. Note that X-ray and VHE data discussed here are not simultaneous. Therefore, all conclusions have to be drawn with caution especially with respect to flaring episodes present in both data sets.

Our proposed blazar classification of IC 310 is further supported by Giommi et al. (2012a). These authors argue that $\sim 90\%$ of moderately beamed blazars in the local universe ($z < 0.07$) are misclassified as radio-galaxies due to their unusually low AGN luminosities: their jets do make small angles with the line of sight to the observer, but their non-thermal emission still gets swamped by the (thermal) light from host galaxy. IC 310 might be such a case.

According to the blazar sequence³ (Fossati et al. 1998; Ghisellini et al. 1998), the apparent luminosity shown in the SED (Fig. 5) would suggest that IC 310 is an extreme high frequency peaked BL Lac object (HBL) with the synchrotron hump peaking in hard X-ray and the Inverse Compton hump peaking in the

multi-TeV regime. IC 310 could be the first detected member of a population of very extreme HBLs that went so far undetected because of their very low luminosity. If so, the accretion rate of the active nucleus of IC 310 should be lower than for typical HBLs, hence of lower luminosity. This explanation is in line with the implications for the cosmic evolution of blazars as discussed by Cavaliere & D’Elia (2002) or Ghisellini & Tavecchio (2008): the classification of different blazar populations might be connected to the fact that we are looking at the same type of object (AGN with small angle to the line of sight) but at different stages of development. For example, blazars with high accretion rate, seen as flat spectrum radio quasars with low mass black holes ($M \approx 10^6 M_{\odot}$) are at the beginning of the life of the AGN and in comparison, BL Lac objects have already accreted most of the material in their environment, resulting in large masses of the black hole ($M > 10^9 M_{\odot}$) and low accretion rate. In this context, the active nucleus of IC 310 with a mass of the black hole of $M_{\text{BH}} > 10^8 M_{\odot}$ (Aleksić et al. 2010) might have already accreted most of the surrounding matter, explaining the relatively high black hole mass, low accretion rate and the low luminosity.

Another possibility is that IC 310 is a slightly misaligned, extreme HBL. In some extreme HBL such as 1ES 1426+428 (Wolter et al. 2008), the synchrotron radiation can peak around 100 keV. Even with a misaligned jet, the observed SED of such objects could peak in the hard X-ray regime. A low apparent luminosity as seen in case of IC 310 could be a consequence of a relatively large angle between the line of sight and the jet. In principle, a factor of ~ 100 in the luminosity can be achieved

³ Recent studies (e.g. Giommi et al. 2012a, 2012b) suggest that the blazar sequence might be affected, or even triggered, by selection effects.

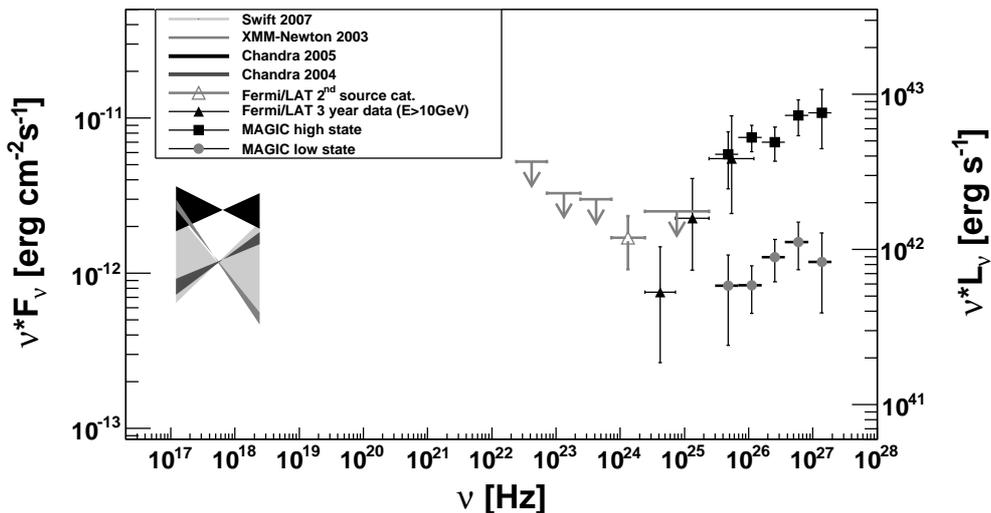


Fig. 5. Spectral energy distribution from X-rays to VHE γ rays. In X-ray, butterflies of the observations are shown. The gray, open triangle and the gray upper limits are obtained from the *Fermi*-LAT second source catalog (Nolan et al. 2012). Filled triangles depict the results from the dedicated high-energy analysis reported here. The MAGIC results (EBL corrected) for the high and low states are marked as full squares and full circles, respectively. The corresponding apparent luminosity is given at the right axis. For the calculation, the luminosity distance of 76.65 Mpc obtained from the NASA Extragalactic Database (NED) has been used.

with an angle of $\sim 8.5^\circ$ (assuming a Lorentz factor $\Gamma = 15$, i.e., a Doppler factor $\delta \sim 5$), consistent with the upper limit from Kadler et al. (2012). TeV flux variability presented here can still be explained within this model, considering an emission size of about 10^{16} cm. Instead, variability on shorter timescales (intra-night) would be challenging for this scenario. The existence of such moderately beamed blazars is expected by Ghisellini & Tavecchio (2008), who predict a large population of weak blazars with angle to the line of sight $\theta = 4 - 7^\circ$. For IC 310, this scenario is supported by the fact that the structure of the jet extends up to ~ 400 kpc (Sijbring & de Bruyn 1998). Such a long extension suggests that the jet kinetic power is not too low. In contrast, the jets of HBLs are very short and basically lose their structure after a few kiloparsecs (Rector et al. 2003).

The flux variability of IC 310 reported here is not faster than the variability of other TeV radio galaxies, e.g. M87 (Aharonian et al. 2006; Abramowski et al. 2012) or NGC 1275 (Colin et al. 2012), which have relatively large angles between the jet and the line of sight. However, non-blazar scenarios developed for M87 such as pulsar-type processes in the central black hole magnetosphere (Neronov & Aharonian 2007) may have difficulties to explain the luminosity of IC 310 with a smaller mass of the central black hole.

5. Summary and Conclusions

In this paper we presented the re-analysis of the IC 310 data taken with the MAGIC telescopes between 2009 October and 2010 February. Using an improved analysis taking into account the data from both wobble positions, we revealed a day-scale flux variability for this object. A high and a low state spectrum of the source were defined and investigated independently. The photon indices of both were comparable within the error bars whereas the flux at 1 TeV is ~ 7 times higher in the high state compared to the low state. Also, we derived a spectrum from *Fermi*-LAT data above 10 GeV of the first three years of operation (2008–2011). Additionally, we studied the source in the X-ray regime by analyzing archival data from *Chandra*, *XMM-*

Newton and *Swift*-XRT. We found clear evidence for flux variability as well as spectral variability on timescales of years.

The non-simultaneous broadband SED appears relatively hard in both X-ray ($\Gamma = 1.76 - 2.55$) and γ ray ($\Gamma = 1.3 - 1.85^4$) energy band. The day-scale VHE variability rules out emission models occurring in a bow shock between the jet and the ICM, but strongly support the blazar-like scenario. In the typical Synchrotron Self-Compton model, the SED of IC 310 can be interpreted as an extreme HBL with synchrotron radiation peaking in X-ray and inverse Compton peaking in the multi-TeV band. However, the intrinsic luminosity is a few orders of magnitude lower than for known extreme HBL. We suggest IC 310 can be either an extreme case of the blazar sequence with a very low accretion rate or a slightly misaligned version of an extreme HBL ($\theta \simeq 10^\circ$). The kpc-scale radio structure, unusual for a BL Lac, supports the misaligned scenario as it suggests a not too weak jet power. Thus, IC 310 could be a member of a transition population between BL Lacs (classical VHE emitters) and FR I radio galaxies.

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⁴ The photon index of 1.85 given here is obtained after the correction for the EBL absorption using the model by Dominguez et al. (2011).

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